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Bounding Lorentz Violation at Particle Colliders By Tracking the Motion of Charged Particles

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Abstract

In the presence of Lorentz violation, the motion of a charged particle in a magnetic field is distorted. By measuring the eccentricities of particles' elliptical orbits and studying how those eccentricities vary with the absolute orientation of the laboratory, it is possible to constrain the Lorentz-violating c_{JK} parameters. For each observed species, this method can provide constraints on four linear combinations of coefficients for which, in some species, there are presently no two-sided bounds.

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1 Introduction

Recently, there has been a surge of interest in the possibility that some symmetries (such as Lorentz and CPT invariances), while apparently fundamental, might not actually hold exactly. Despite a large number of experimental searches, no compelling evidence for violation of Lorentz invariance or CPT has been found. However, new Lorentz and CPT tests are still underway, and the study of these symmetries remains an active area of research—both experimentally and theoretically.

If violations of Lorentz invariance or CPT are uncovered, that would be a discovery of the utmost importance. Such symmetry violations would be immediate evidence for new physics beyond the standard model with a completely new structure. The mere existence of such symmetry breaking would tell us a tremendous amount about the character of fundamental physics. For example, a breaking of Lorentz symmetry could be tied to the structure of quantum gravity.

Unlike some possible forms of exotic new physics that have been suggested, Lorentz and CPT violations can be studied in the framework of quantum field theory. Violations of these symmetries may be described in an effective field theory called the standard model extension (SME). The SME contains Lorentz- and CPT-violating corrections to the standard model, parameterized by small tensor-valued background fields [1, 2]. This field theory framework can also be expanded to cover gravity [3]. Both the renormalizability [4, 5] and stability [6] of the SME have been studied. The minimal SME (which contains only gauge-invariant, renormalizable forms of Lorentz violation) has become the standard framework used for parameterizing the results of experimental Lorentz tests.

Recent searches for Lorentz violation have included studies of matter-antimatter asymmetries for trapped charged particles [7, 8, 9] and bound state systems [10, 11], measurements of muon properties [12, 13], analyses of the behavior of spin-polarized matter [14], frequency standard comparisons [15, 16, 17, 18], Michelson-Morley experiments with cryogenic resonators [19, 20, 21, 22, 23], Doppler effect measurements [24, 25], measurements of neutral meson oscillations [26, 27, 28, 29, 30, 31], polarization measurements on the light from cosmological sources [32, 33, 34, 35], high-energy astrophysical tests [36, 37, 38, 39, 40], precision tests of gravity [41, 42], and others. The results of these experiments set constraints on the various SME coefficients, and up-to-date information about most of these constraints may be found in [43].

There are many precise bounds on the forms of Lorentz violation that affect stable particles, but Lorentz violation with unstable particles is much more difficult to study. There have been some impressive measurements made using muon spin precession, neutral meson oscillations, pion decay [44], and neutrons (both free neutrons [45] and bound ones in stable nuclei). However, there are still relatively few constraints on the SME parameters for most second- and third-generation fields. Unstable particles are produced and detected in copious numbers in collider experiments, and so it is natural to look for ways to test Lorentz symmetry using new or existing collider data sets. (Data from accelerators has

already been used to place bounds on electron and photon Lorentz violations [46, 47, 48].)

Collider data that is considered to have little value in searches for novel (but still Lorentz-invariant) physics may actually be fairly sensitive to violations of Lorentz symmetry. In the paper, we shall identify an observable that may be used to test Lorentz invariance for unstable charged species. In magnetized detectors, ultrarelativistic particles trace out synchrotron orbits, and the shapes of these orbits are affected by the SME coefficients. We shall show how an analysis of this effect can be used to place new constraints on SME parameters.

Section 2 of this paper will introduce the form of Lorentz violation that is most relevant for particle tracking experiments. The shapes of synchrotron orbits in the presence of this kind of Lorentz violation have already been calculated [49], and it shall be shown that the modified motion gives rise to orbits with nonvanishing eccentricities. The analysis of these orbits continues in section 3, which describes how sidereal variations in orbital eccentricities can be related to the underlying SME parameters. This section is subdivided into two cases: section 3.1 considers the case in which Lorentz violation is the dominant source of orbital deformation; and section 3.2 deals with the possibility that other systematic effects may largely determine an orbit’s eccentricity, with Lorentz violation as an additional small perturbation. Some comments and conclusions are given in section 4.

2 Lorentz-Violating Synchrotron Motion

We shall be studying the quantum electrodynamics (QED) sector of the SME. The QED Lagrange density containing the form of Lorentz violation that is most relevant at high energies is

$$\mathcal{L} = -\frac{1}{4}F^{\mu\nu}F_{\mu\nu} + \bar{\psi}[(\gamma^\mu + c^{\nu\mu}\gamma_\nu)(i\partial_\mu - qA_\mu) - m]\psi. \quad (1)$$

Spin-independent Lorentz violation at high energies is controlled primarily by the $c^{\nu\mu}$ coefficients. There is a separate set of these coefficients for each fermion species. The importance of the c coefficients grows at high energies. There is another, similar set of coefficients—the $d^{\nu\mu}$ —which also become more important at high energies, but the effects of the d parameters depend on helicity and average to zero for unpolarized particles. There is also the potential for Lorentz violation in the electromagnetic sector; however, photon Lorentz violation can be constrained sufficiently accurately in low-energy, high precision experiments that it may be neglected in this analysis of accelerator-based Lorentz tests.

At leading order, $c^{\nu\mu}$ may be taken to be traceless and symmetric, so it contains nine independent parameters for each species. The energy of a particle, including the leading order effects of c is

$$E = \sqrt{(m^2 + \vec{\pi}^2)(1 - 2c_{00}) - 2c_{jk}\pi_j\pi_k - 2c_{0j}\pi_j}, \quad (2)$$

where $\vec{\pi}$ is the mechanical momentum $\vec{p} - q\vec{A}$. The corresponding velocity is [50]

$$v_j = \frac{\pi_j}{\sqrt{m^2 + \vec{\pi}^2}} - c_{00} \frac{\pi_j}{\sqrt{m^2 + \vec{\pi}^2}} - 2c_{jk} \frac{\pi_k}{\sqrt{m^2 + \vec{\pi}^2}} + c_{kl} \frac{\pi_j \pi_k \pi_l}{(m^2 + \vec{\pi}^2)^{3/2}} - 2c_{0j}. \quad (3)$$

For stable charged particles, there are many possible ways to constrain the c coefficients. However, for unstable particles, there are very few bounds; in most cases, essentially the only bounds come from the absence of the photon decay process $\gamma \rightarrow X^+ + X^-$, which may be allowed for initial photon energies $E \gtrsim m_X / \sqrt{|c_X|}$. Here c_X represents a linear combination of the c coefficients for the particle species X ; precisely which linear combination can be constrained with a given observation depends on the direction of the photons involved, although the coefficients are generally of $\mathcal{O}(1)$. However, the photon decay bounds get worse for heavier particles, since the threshold energy increases with the particle mass m_X ; bounds based on the survival of 10–100 TeV photons over astrophysical distances are typically at a level $\sim 10^{-9} [m_X / (1 \text{ GeV})]^2$. Moreover, these are one-sided bounds, and a sufficiently large value of c_{00} (the coefficient of pure boost invariance violation) could hide the effects of any other nonzero c coefficients.

This means that any two-sided constraints on the c coefficients for unstable species will exclude new regions of the SME parameter space. Constraints comparable to (or stronger than) the photon survival bounds would be especially interesting. For muons, achieving this degree of precision may be a challenge, but for heavier species, the astrophysical bounds grow weaker quite rapidly. Accelerator constraints may be competitive, since they would not be expected to depend so strongly on the particle mass.

In this paper, we shall be concerned with the c_{jk} coefficients only—specifically, the five linear combinations of these coefficients that break rotational isotropy. The orbits of relativistic charged particles in a magnetic field are not as sensitive to other forms of Lorentz violation. For example, the d coefficients have no net affect on particle orbits when there are equal populations of positive and negative helicity fermions. The parameter c_{00} breaks boost invariance only, and so it cannot be constrained by measurements of spatial isotropy. Moreover, the spatial trace c_{jj} is equivalent to c_{00} by the (four-dimensional) tracelessness of $c^{\nu\mu}$.

The trickiest parameters to rule out of consideration are the boost anisotropy parameters c_{0j} . These are odd under parity (P) and time reversal (T), and they do not affect the ellipticities of complete orbits. Since c_{0j} provides merely a constant offset to the velocity of a particle, at nonrelativistic energies where no particles are created or destroyed, c_{0j} is completely unobservable. This is a consequence of the Galilean invariance of nonrelativistic physics. While c_{0j} does have physical effects upon relativistic experiments, these coefficients will still only affect parity-odd observables.

In fact, if energy losses are neglected, the effect of the c_{0j} coefficients alone is to shift the velocity to $v_j = \pi_j / \sqrt{m^2 + \vec{\pi}^2} - c_{0j}$. The equation of motion in a magnetic field, $\frac{d\vec{\pi}}{dt} = q(\vec{v} \times \vec{B})$ is not changed by the Lorentz violation (as a consequence of gauge invariance). So $\frac{d\vec{v}}{dt} = \frac{q}{\gamma m}(\vec{v} \times \vec{B})$ continues to hold, even in the presence of the c_{0j} . In

other words, when synchrotron losses are neglected, c_{0j} has no effect on a particle's orbit. If energy losses are included, $\frac{dv_j}{dt}$ does acquire a new term $(\dot{E}/E)c_{0j}$, in addition to the terms ordinarily present. This is potentially observable, but its effects are suppressed by the energy loss rate, and the novel force does not affect the eccentricity of adiabatically shrinking orbits.

We shall therefore neglect the c_{00} and c_{0j} , restricting our attention to c_{jk} Lorentz violation only. The motion of a particle with this kind of Lorentz violation in a constant background magnetic field \vec{B} can be calculated exactly, if the radiation by the accelerated charges is neglected. To first order in the c_{jk} coefficients, the velocity of an orbiting charge is [49]

$$\begin{bmatrix} v_1(t) \\ v_2(t) \\ v_3(t) \end{bmatrix} = \begin{bmatrix} v_{10} (\cos \omega t + 2c_{12} \sin \omega t) \\ -v_{10}(1 + c_{11} - c_{22}) \sin \omega t \\ 2v_{10} [c_{23} \sin \omega t - c_{13} (\cos \omega t - 1)] + v_{30} \end{bmatrix}. \quad (4)$$

The direction \hat{B} of the magnetic field is taken to be along the z -axis, and the initial conditions are that $\vec{v}(0) = (v_{10}, 0, v_{30})$. The frequency $\omega = (1 - c_{11} - c_{22})(eB/E)$ differs from the usual expression by small Lorentz-violating corrections, although the energy $E = \sqrt{m^2 + (\delta_{jk} - 2c_{jk})\pi_j\pi_k}$ remains a constant of the motion.

As in conventional synchrotron motion, there are two components to the charged particle's trajectory—periodic motion superposed with a constant velocity along the direction \hat{B} of the field. The periodic elliptical motion is not precisely restricted to the plane perpendicular to \vec{B} , but the tilt out of this plane is small. The projection of the orbit in the plane normal to \vec{B} is an ellipse, which generally has nonzero eccentricity; this eccentricity is the key indicator that rotation symmetry does not hold exactly. When radiation emission is included, the size of an elliptical track will decrease adiabatically if the energy losses are gradual enough. The ellipse will shrink slowly, but its eccentricity and orientation will remain unchanged.

The shape of the orbit in the directions perpendicular to \vec{B} is governed by the projection of c_{jk} onto those two directions:

$$c_{jk}^\perp = c_{jk} - 2(c_{jl}\hat{B}_l)\hat{B}_k + (c_{lm}\hat{B}_l\hat{B}_m)\hat{B}_j\hat{B}_k. \quad (5)$$

It is evident from setting $c_{12} = 0$ in (4) that the charge's oscillatory motions along the the eigenvectors of c^\perp (neglecting the eigenvector \hat{B} with vanishing eigenvalue) are out of phase by $\frac{\pi}{2}$. In fact, when the orbital ellipse is projected into the plane normal to \hat{B} , the major and minor axes are oriented along those eigenvectors of c^\perp . Calling unit vectors in these directions \hat{a} for the major axis and \hat{b} for the minor, the ratio of the two axes is

$$\frac{a}{b} = 1 - c_{jk}^\perp \hat{a}_j \hat{a}_k + c_{jk}^\perp \hat{b}_j \hat{b}_k. \quad (6)$$

$c_{jk}^\perp \hat{a}_j \hat{a}_k$ and $c_{jk}^\perp \hat{b}_j \hat{b}_k$ are the two (potentially nonzero) eigenvalues of c^\perp . The eigenvector of c^\perp corresponding to the smaller eigenvalue indicates the major axis direction; the larger eigenvalue indicates the minor axis.

So the orbital eccentricity squared is $\epsilon^2 = 2|\lambda_1 - \lambda_2|$ —proportional to the difference between the larger and smaller eigenvalues of the 2×2 matrix formed by projecting c_{jk} into the plane normal to \vec{B} . Henceforth, we shall refer to this 2×2 matrix as c^\perp , excising the row and column of zeros the correspond to the \hat{B} -direction. The matrix may be constructed explicitly by choosing two unit vector \hat{e}_1 and \hat{e}_2 , so that \hat{e}_1 , \hat{e}_2 , and \hat{B} form an orthonormal basis. Then the 2×2 projection of c_{jk} is

$$c^\perp = \begin{bmatrix} c_{jk}\hat{e}_{1j}\hat{e}_{1k} & c_{jk}\hat{e}_{1j}\hat{e}_{2k} \\ c_{jk}\hat{e}_{2j}\hat{e}_{1k} & c_{jk}\hat{e}_{2j}\hat{e}_{2k} \end{bmatrix}. \quad (7)$$

The eigenvalue difference is then $(\lambda_1 - \lambda_2) = \sqrt{(c_{11}^\perp - c_{22}^\perp)^2 + 4(c_{12}^\perp)^2}$ in terms of the elements of c^\perp .

3 Experimental Observables

The eccentricity of a synchrotron orbit is an experimental observable, which can be related to the underlying c_{jk} coefficients for the revolving species. However, the relationship between ϵ and the c_{jk} in section 2 was derived under the assumption that Lorentz violation was the largest effect driving the orbits to have systematic ellipticities. The experimental observables in such a regime are derived in section 3.1. However, if there are systematic effects that dominate the orbital eccentricity, the observables are different and somewhat simpler; they are explained in section 3.2. Which of the two eccentricity regimes is physically relevant must be inferred from observational data.

The same subset of the c_{jk} coefficients can be constrained in either case. Experimental constraints on the SME parameters are conventionally expressed in a system of Sun-centered celestial equatorial coordinates [51]. Coordinates in this system are denoted by capital letters. From the origin at the center of the Sun, the Z -axis points along the Earth's rotation axis; the X -axis points toward the vernal equinox point on the celestial sphere; and the Y -axis is determined by the right hand rule.

3.1 Eccentricities Dominated by Lorentz Violation

If Lorentz violation is the largest source of orbital eccentricity, ϵ can be expressed in terms of the c_{JK} parameters in the Sun-centered frame and the components of \hat{B} . The explicit form of the eigenvalue difference squared, $(\lambda_1 - \lambda_2)^2$, is

$$\begin{aligned} (\lambda_1 - \lambda_2)^2 = & \sum_{J=1}^3 (1 - \hat{B}_J^2)^2 c_{JJ}^2 + \sum_{J \neq K} \left\{ -4\hat{B}_J\hat{B}_K(1 - \hat{B}_J^2)c_{JJ}c_{JK} \right. \\ & \left. + [2\hat{B}_J^2\hat{B}_K^2 - (1 - \hat{B}_J^2)(1 - \hat{B}_K^2)]c_{JJ}c_{KK} + 2(1 - \hat{B}_J^2)(1 - \hat{B}_K^2)c_{JK}^2 \right\} \\ & + \sum_{J \neq K \neq L \neq J} \left[4\hat{B}_K\hat{B}_L(1 + \hat{B}_J^2)c_{JJ}c_{KL} - 4\hat{B}_J\hat{B}_K(1 - \hat{B}_L^2)c_{JL}c_{KL} \right]. \quad (8) \end{aligned}$$

The sums run over all combinations of distinct indices, and there are no additional sums implied by repeated indices. Note that this means that each of the $c_{JJ}c_{KK}$, c_{JK}^2 , and $c_{JL}c_{KL}$ terms appears twice, with J and K exchanged; the total coefficient for a given one of these terms is thus twice what what might be naively assumed from a casual look at (8).

At any given instant, the magnetic field provides a fixed axis in space $\vec{B} = B\hat{B}$. This axis rotates with the Earth. If \vec{B} lies in the plane tangent to the Earth's surface at the location of the lab (at colatitude χ), then the time-dependent unit vector $\hat{B}(t)$ is

$$\begin{aligned}\hat{B}(t) &= -(\cos\psi\cos\Omega t + \sin\psi\cos\chi\sin\Omega t)\hat{X} + \\ &\quad (\cos\psi\sin\Omega t - \sin\psi\cos\chi\cos\Omega t)\hat{Y} + \sin\psi\sin\chi\hat{Z}\end{aligned}\tag{9}$$

$$\equiv (\alpha\cos\Omega t + \beta\sin\Omega t)\hat{X} + (\beta\cos\Omega t - \alpha\sin\Omega t)\hat{Y} + \gamma\hat{Z}.\tag{10}$$

The angle ψ denotes the orientation of \vec{B} in the lab frame; ψ is the angle of \vec{B} north of the eastward direction. Ω is the Earth's sidereal rotation frequency, and t is local sidereal time. The orientation of \vec{B} depends on the three coefficients $\alpha = -\cos\psi$, $\beta = -\sin\psi\cos\chi$ and $\gamma = \sin\psi\sin\chi$, with $\alpha^2 + \beta^2 + \gamma^2 = 1$. α , β , and γ are determined by the specific geometry of an individual collider experiment, as they parameterize the direction of the local magnetic field in the laboratory.

The \hat{B}_X and \hat{B}_Y components of the field are time dependent, but \hat{B}_Z is not. According to (8), the coefficients of the $c_{JK}c_{LM}$ in $(\lambda_1 - \lambda_2)^2$ are polynomials in \hat{B}_X and \hat{B}_Y with degrees of up to four. This means that $(\lambda_1 - \lambda_2)^2$ can be expressed in terms of Fourier components as

$$\begin{aligned}(\lambda_1 - \lambda_2)^2 &= \mathcal{A}_0 + \mathcal{A}_\Omega\cos\Omega t + \mathcal{B}_\Omega\sin\Omega t + \mathcal{A}_{2\Omega}\cos 2\Omega t + \mathcal{B}_{2\Omega}\sin 2\Omega t \\ &\quad + \mathcal{A}_{3\Omega}\cos 3\Omega t + \mathcal{B}_{3\Omega}\sin 3\Omega t + \mathcal{A}_{4\Omega}\cos 4\Omega t + \mathcal{B}_{4\Omega}\sin 4\Omega t.\end{aligned}\tag{11}$$

The Fourier coefficients \mathcal{A} and \mathcal{B} , expressed in terms of the c_{JK} , are rather complicated. However, they can be simplified substantially with knowledge of which linear combinations of the coefficients this technique is sensitive to.

In the Fourier series (11), the only coefficient that depends on c_{ZZ}^2 is \mathcal{A}_0 . The reason is that the the orientation of the magnetic field relative to the Z -axis does not change as the Earth rotates. This means that there is no sensitivity to the specific linear combination of coefficients $c_Q = c_{XX} + c_{YY} - 2c_{ZZ}$ (which parameterizes the effects of a preferred direction parallel to the Z -axis) unless at least one other linear combination of c_{JK} coefficients is nonzero. [The formula for \mathcal{A}_0 does contain a term $\frac{1}{4}(\alpha^2 + \beta^2)c_Q^2$. However, while \mathcal{A}_0 represents a genuine signal for anisotropy, the lack of any sidereal dependence tends to make this signal indistinguishable from the effects of laboratory systematics.]

Most experiments that rely on the Earth's rotation do not exhibit any sensitivity to c_Q at $\mathcal{O}(c)$. However, since the observable a/b is a non-analytic function of the c_{JK} , the situation here differs a bit from the usual one. There is no sensitivity to c_Q alone.

However, the sidereal oscillations in a/b do depend on the geometric means of c_Q and other linear combinations of the Lorentz violation coefficients. So unless there is Lorentz violation of another sort, measurements of the \mathcal{A} and \mathcal{B} coefficients cannot provide any constraints on c_Q . In order to avoid ill conditioning in the problem of extracting the c_{JK} coefficients from measured values of \mathcal{A} and \mathcal{B} , it makes sense to set $c_Q = 0$. Doing so will simplify the expressions for the \mathcal{A} and \mathcal{B} .

Moreover, since there is no sensitivity to the trace of c_{JK} (which describes only boost invariance violation), c_{XX} and c_{YY} can only enter in the combination $c_- = c_{XX} - c_{YY}$. The forms of the coefficients of the time-dependent Fourier components of are

$$\begin{aligned} \mathcal{A}_\Omega &= \beta\gamma(\alpha^2 + \beta^2 - 4)c_-c_{XZ} + \alpha\gamma(\alpha^2 + \beta^2 - 4)c_-c_{YZ} \\ &\quad - 2\alpha\gamma(\alpha^2 + \beta^2 - 4)c_{XY}c_{XZ} + 2\beta\gamma(\alpha^2 + \beta^2 - 4)c_{XY}c_{YZ} \\ &\quad [+2\beta\gamma(\alpha^2 + \beta^2)c_Qc_{XZ} - 2\alpha\gamma(\alpha^2 + \beta^2)c_Qc_{YZ}] \end{aligned} \quad (12)$$

$$\begin{aligned} \mathcal{B}_\Omega &= \alpha\gamma(\alpha^2 + \beta^2 - 4)c_-c_{XZ} - \beta\gamma(\alpha^2 + \beta^2 - 4)c_-c_{YZ} \\ &\quad + 2\beta\gamma(\alpha^2 + \beta^2 - 4)c_{XY}c_{XZ} + 2\alpha\gamma(\alpha^2 + \beta^2 - 4)c_{XY}c_{YZ} \\ &\quad [+2\alpha\gamma(\alpha^2 + \beta^2)c_Qc_{XZ} + 2\beta\gamma(\alpha^2 + \beta^2)c_Qc_{YZ}] \end{aligned} \quad (13)$$

$$\begin{aligned} \mathcal{A}_{2\Omega} &= +2(\alpha^4 - \beta^4)c_{XZ}^2 + 2(\beta^4 - \alpha^4)c_{YZ}^2 + 8\alpha\beta(\alpha^2 + \beta^2)c_{XZ}c_{YZ} \\ &\quad [-\frac{1}{2}(\alpha^2 - \beta^2)(\alpha^2 + \beta^2 - 2)c_-c_Q - 2\alpha\beta(\alpha^2 + \beta^2 - 2)c_Qc_{XY}] \end{aligned} \quad (14)$$

$$\begin{aligned} \mathcal{B}_{2\Omega} &= -4\alpha\beta(\alpha^2 + \beta^2)c_{XZ}^2 + 4\alpha\beta(\alpha^2 + \beta^2)c_{YZ}^2 + 4(\alpha^4 - \beta^4)c_{XZ}c_{YZ} \\ &\quad [+ \alpha\beta(\alpha^2 + \beta^2 - 2)c_-c_Q - (\alpha^2 - \beta^2)(\alpha^2 + \beta^2 - 2)c_Qc_{XY}] \end{aligned} \quad (15)$$

$$\begin{aligned} \mathcal{A}_{3\Omega} &= \beta\gamma(\beta^2 - 3\alpha^2)c_-c_{XZ} + \alpha\gamma(\alpha^2 - 3\beta^2)c_-c_{YZ} \\ &\quad + 2\alpha\gamma(\alpha^2 - 3\beta^2)c_{XY}c_{XZ} + 2\beta\gamma(3\alpha^2 - \beta^2)c_{XY}c_{YZ} \end{aligned} \quad (16)$$

$$\begin{aligned} \mathcal{B}_{3\Omega} &= \alpha\gamma(3\beta^2 - \alpha^2)c_-c_{XZ} + \beta\gamma(\beta^2 - 3\alpha^2)c_-c_{YZ} \\ &\quad + 2\beta\gamma(\beta^2 - 3\alpha^2)c_{XY}c_{XZ} + 2\alpha\gamma(\alpha^2 - 3\beta^2)c_{XY}c_{YZ} \end{aligned} \quad (17)$$

$$\mathcal{A}_{4\Omega} = \frac{1}{8}(\alpha^4 - 6\alpha^2\beta^2 + \beta^4)c_-^2 + 2\alpha\beta(\alpha^2 - \beta^2)c_-c_{XY} - \frac{1}{2}(\alpha^4 - 6\alpha^2\beta^2 + \beta^4)c_{XY}^2 \quad (18)$$

$$\mathcal{B}_{4\Omega} = \frac{1}{2}\alpha\beta(\beta^2 - \alpha^2)c_-^2 + \frac{1}{2}(\alpha^4 - 6\alpha^2\beta^2 + \beta^4)c_-c_{XY} + 2\alpha\beta(\alpha^2 - \beta^2)c_{XY}^2. \quad (19)$$

The expressions are given in terms of c_- , c_{XY} , c_{XZ} , c_{YZ} , and c_Q . With the exception of c_Q , only one of these linear combinations needs to be nonzero for there to be an observable effect. The terms that depend on c_Q are provided for the sake of completeness only; these terms are printed in square brackets, so that they can be dropped in practical calculations.

The expressions for the \mathcal{A} and \mathcal{B} Fourier coefficients are still complicated, but they possess a number of readily understandable features. For example, if $\gamma = 0$, indicating that \vec{B} has no component in the Z -direction, the Fourier coefficients for odd harmonics all vanish. In this case, the plane normal to \vec{B} is parallel to the Z -axis. As the laboratory rotates around that axis, the plane returns to its original orientation after only half a

sidereal day. However, if $\gamma \neq 0$, the normal plane is tilted relative to the Z -axis, and it takes a full day for it to return to its original orientation.

If the Lorentz violation is purely in the XY -plane (only $c_{XX} = -c_{YY}$ and c_{XY} nonzero) only $\mathcal{A}_{4\Omega}$ and $\mathcal{B}_{4\Omega}$ are nonvanishing. The major and minor axes of the orbit's projection interchange positions (in the lab frame) every quarter sidereal day; the ratio of the major to the minor thus oscillates four times daily. Moreover, the amplitude of the oscillations in $(\lambda_1 - \lambda_2)^2$ is proportional to $(c_{XX} - c_{YY})^2 + 4c_{XY}^2$, which in this case is the square of the difference of the nonzero eigenvalues of the matrix c_{JK} itself.

To place constraints on the c_{JK} coefficients requires a fit of the observed $\epsilon^4 = 4(a/b-1)^2$ data for a given species. The fit should identify oscillations at frequencies Ω , 2Ω , 3Ω , and 4Ω . Fermion and antifermion data can be combined and fitted together, because the c coefficients distort particle and antiparticle orbits in the same fashion. The formulas for the Fourier coefficients given in (12–19) can then be inverted through a second fit, giving experimental values of the c_{JK} .

There are eight potentially nonzero Fourier coefficients in (11), making the c_{JK} values strongly overconstrained. Moreover, there are additional observables in this system, which depend on the same underlying coefficients. If a significant sidereal oscillation signal is seen in the ratio a/b , the potential Lorentz violation can be further checked by looking at the actual directions of the major and minor axes, which will also oscillate in sidereal time.

3.2 Eccentricities Dominated by Systematics

Section 3.1 discussed the behavior of the eccentricity in an idealized situation, in which Lorentz violation is the dominant effect that deforms the synchrotron orbits away from circularity. This is a potentially important regime, and the results are also interesting theoretically, since they introduce an $\mathcal{O}(c)$ observable that is a non-analytic function of the c_{JK} parameters. However, other effects may also lead to non-circular orbits, and a different analysis is required if those other effects are more important than any Lorentz violation.

It would be conceivable to have a systematic effect that made all orbits elliptical, with the axis directions \hat{a} and \hat{b} fixed in the laboratory frame. Alternatively, systematics might lead to a nonzero average eccentricity, but without fixed \hat{a} and \hat{b} . This might arise from changes in the detector composition as a function of the distance from the beam axis. Different materials, in which the orbiting particles have different energy loss rates, could lead to an orbital deformation for which \hat{a} or \hat{b} is always oriented along the radial direction.

If the orbital eccentricity is dominated by systematics, the effects of the c_{JK} may be treated as a perturbation. The primary Lorentz-violating effect is still an extension of the major axis by a factor $1 - c_{jk}^\perp \hat{a}_j \hat{a}_k = 1 - c_{jk} \hat{a}_j \hat{a}_k$ and contraction of the minor axis by $1 + c_{jk} \hat{b}_j \hat{b}_k$. (Since \hat{a} and \hat{b} are then defined to lie in the plane normal to \vec{B} , the use of c^\perp

instead of c is unnecessary.) The difference from the case discussed previously is that \hat{a} and \hat{b} are not determined by c itself but by the systematics. This makes the eccentricity

$$\epsilon^2 = \epsilon_0^2 - 2(1 - \epsilon_0^2)(c_{jk}\hat{a}_j\hat{a}_k - c_{jk}\hat{b}_j\hat{b}_k), \quad (20)$$

where ϵ_0 is the eccentricity caused by systematics unrelated to the Lorentz violation.

Since ϵ_0 should not exhibit sidereal variations, it is still possible to constrain the c_{JK} coefficients by looking at the time dependence of ϵ . However, the data will need to be binned according to the direction \hat{a} . If \hat{a} is always radial and charged particles are emitted from the interaction region in all directions, then the average contribution that the c_{JK} make to ϵ^2 vanishes. For some orbital orientations, $-2(1 - \epsilon_0^2)(c_{jk}\hat{a}_j\hat{a}_k - c_{jk}\hat{b}_j\hat{b}_k)$ will increase ϵ^2 , but for an equal fraction of orbits, the contribution to ϵ^2 will be negative. However, if different \hat{a} values are analyzed separately, this cancellation does not occur.

For orbits with major axis direction \hat{a} , the minor axis lies along $\hat{b} = \hat{a} \times \hat{B}$. If $\hat{a}(0) = \delta\hat{X} + \zeta\hat{Y} + \eta\hat{Z}$ at $t = 0$, then the time-dependent axis directions are

$$\hat{a}(t) = (\delta \cos \Omega t + \zeta \sin \Omega t)\hat{X} + (\zeta \cos \Omega t - \delta \sin \Omega t)\hat{Y} + \eta\hat{Z} \quad (21)$$

$$\begin{aligned} \hat{b}(t) = & [(\gamma\zeta - \beta\eta) \cos \Omega t + (\alpha\eta - \gamma\delta) \sin \Omega t]\hat{X} \\ & + [(\alpha\eta - \gamma\delta) \cos \Omega t - (\gamma\zeta - \beta\eta) \sin \Omega t]\hat{Y} + (\beta\delta - \alpha\zeta)\hat{Z}. \end{aligned} \quad (22)$$

This makes the observable quantity ϵ^2 a second-order polynomial in $\cos \Omega t$ and $\sin \Omega t$, expandable as

$$\epsilon^2 = -2(1 - \epsilon_0^2) (\mathcal{C}_0 + \mathcal{C}_\Omega \cos \Omega t + \mathcal{D}_\Omega \sin \Omega t + \mathcal{C}_{2\Omega} \cos 2\Omega t + \mathcal{D}_{2\Omega} \sin 2\Omega t). \quad (23)$$

The \mathcal{C} and \mathcal{D} Fourier coefficients that describe the sidereal variations in this observable are

$$\mathcal{C}_\Omega = 2[\delta\eta - (\gamma\zeta - \beta\eta)(\beta\delta - \alpha\zeta)]c_{XZ} + 2[\zeta\eta - (\alpha\eta - \gamma\delta)(\beta\delta - \alpha\zeta)]c_{YZ} \quad (24)$$

$$\mathcal{D}_\Omega = 2[\zeta\eta - (\alpha\eta - \gamma\delta)(\beta\delta - \alpha\zeta)]c_{XZ} - 2[\delta\eta - (\gamma\zeta - \beta\eta)(\beta\delta - \alpha\zeta)]c_{YZ} \quad (25)$$

$$\mathcal{C}_{2\Omega} = \frac{1}{2}[\delta^2 - \zeta^2 - (\gamma\zeta - \beta\eta)^2 + (\alpha\eta - \gamma\delta)^2]c_- + 2[\delta\zeta(\gamma\zeta - \beta\eta)(\alpha\eta - \gamma\delta)]c_{XY} \quad (26)$$

$$\mathcal{D}_{2\Omega} = [\delta\zeta(\gamma\zeta - \beta\eta)(\alpha\eta - \gamma\delta)]c_- - [\delta^2 - \zeta^2 - (\gamma\zeta - \beta\eta)^2 + (\alpha\eta - \gamma\delta)^2]c_{XY}. \quad (27)$$

The formulas depend on the directions of \hat{a} and \vec{B} in the laboratory frame.

Generically, the sidereal variations in the eccentricity depend on the same four c_{JK} coefficients that were found to be observable in section 3.1: c_- , c_{XY} , c_{XZ} , and c_{YZ} . Since the effects of the Lorentz violation are, in this case, perturbative—being added onto a larger systematic eccentricity—the \mathcal{C} and \mathcal{D} coefficients are simpler than the \mathcal{A} and \mathcal{B} . The \mathcal{C} and \mathcal{D} depend only linearly on the c_{JK} , rather than quadratically; this is related to the fact that the natural observable is $\epsilon^2 = 2(a/b - 1)$, rather than ϵ^4 . The eight \mathcal{A} and \mathcal{B} Fourier coefficients provide redundant measurements of the four relevant Lorentz

violation parameters, but the \mathcal{C} and \mathcal{D} only yield four such measurements. However, in both regimes there is additional redundancy tied to measurement of the actual direction \hat{a} . In the present analysis, that redundancy is explicit in the fact that the ϵ^2 data needs to be binned according to the laboratory value of \hat{a} , so that each bin provides a separate measurement of the eccentricity.

4 Conclusion

The Lorentz-violating background tensor c_{JK} contains five coefficients parameterizing anisotropy in the fermion sector. One of these, c_Q , can be measured by looking at sidereal variations in the orbital shape only if another linear combination is nonzero and dominates any systematic effects. This leaves four combinations of the c_{JK} to be bounded independently.

For unstable particles, almost any bounds that can be placed on these coefficients would be of interest, since it is quite difficult to constrain Lorentz violation for second- and third-generation species. Measurements of the c coefficients for the t quark, for example, are considered quite interesting, even though they may attain only 10^{-2} – 10^{-1} level sensitivity. However, it would be especially interesting if the laboratory constraints could reach the same level of precision as the astrophysical photon survival bounds. For B mesons, this would mean placing bounds at approximately the 10^{-7} level. Achieving high sensitivities could be assisted by using the large statistics that are available with collider data and by focusing on oscillations at a multiple of the Earth’s sidereal frequency. However, for orbits with $\gtrsim 1$ m dimensions, reaching 10^{-7} precision for the c_{JK} would require determining differences in average orbital dimensions at the sub- μm scale, an extremely challenging proposition. With $100 \mu\text{m}$ precision, two-sided constraints at the 10^{-4} level should be possible, and this would still represent a major accomplishment for third-generation particles.

Although Lorentz violation is often associated with CPT violation, the c coefficients are all even under C and CPT. For this reason, comparisons of the orbital shapes for particles and antiparticles are of rather limited interest. Renormalizable forms of Lorentz violation that do violate CPT tend to diminish in importance at high energies. The d coefficients violate C (but not CPT) and do not become less important at larger energies; however, the effects of the d parameters depend on both particle versus antiparticle identity and on helicity, making them hard to measure without simultaneous measurements of particle polarizations.

While the orbital motion (4) was originally derived for fermionic particles, there is an analogous set of coefficients for charged boson species. The orbits of bosons are affected by Lorentz violation in essentially the same way as the orbits of fermions. To apply the formulas in this paper to bosons, only the replacement $c_{JK} \rightarrow \frac{1}{2}k_{JK}$ is required.

In summary, the Lorentz-violating c coefficients break spatial isotropy. Consequently,

a particle with nonzero c_{jk} coefficients will orbit elliptically (instead of circularly) in a constant magnetic field. In an Earthbound detector, the magnetic field direction will change as the Earth rotates, and so the eccentricity ϵ of the orbital ellipse will be modulated at multiples of the planet's sidereal rotation frequency. Measurements of these modulations can be turned around to extract bounds on four of the five c coefficients that describe spin-independent spatial anisotropy. For most unstable fermion species, these would represent the first two-sided constraints on these important SME parameters.

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